Dark Energy in the Universe

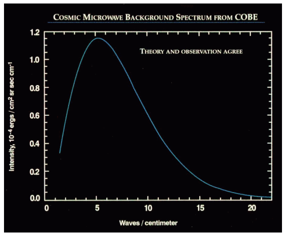
Scott Dodelson

March 25, 2003

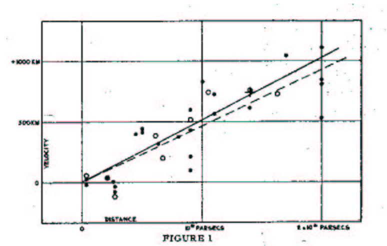
Overview

Standard cosmological model

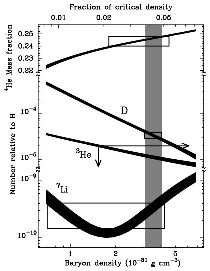
- Cosmic Microwave Background (CMB)



- Hubble expansion



- Light element abundances

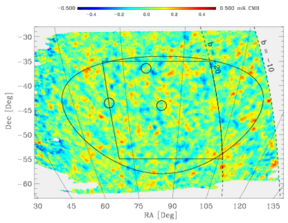


Overview

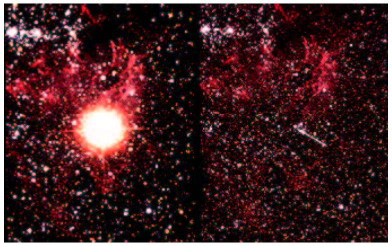
Beyond the standard model

- Dark Matter

- Inflation



- Dark Energy



Overview

Evidence for Dark Energy

- Age: Hubble constant $+$ globular clusters
- Luminosity distance: Type Ia Supernovae
- Cosmic inventory: CMB ($\Omega = 1$) $+$ Many ($\Omega_m \simeq 0.3$)
- Growth function: Weak lensing & Cluster counts

Overview

What is it?

- Cosmological constant Λ: Historical edge (Einstein), very unlikely
- $\Lambda = 0$; transient energy, eventually will go to zero: Modern favorite, very unlikely

Evidence for Dark Energy

Expansion determined by Einstein Equations. If the universe is flat, then

$$\underbrace{\left(\frac{da/dt}{a}\right)^2}_{\text{"kinetic energy"}} = \underbrace{\frac{8\pi G}{3}\rho}_{\text{"potential energy"}}$$

and

$$\frac{d^2a}{dt^2} = -\frac{4\pi G\rho a}{3}\,(1+3w)$$

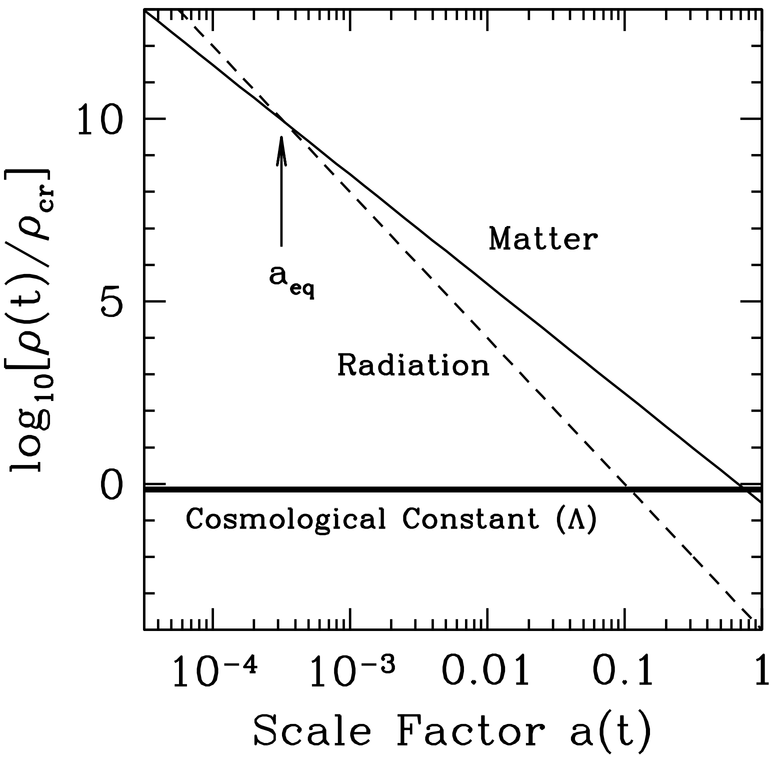
with $w \equiv P/\rho$.

Decceleration unless $w < 0$.

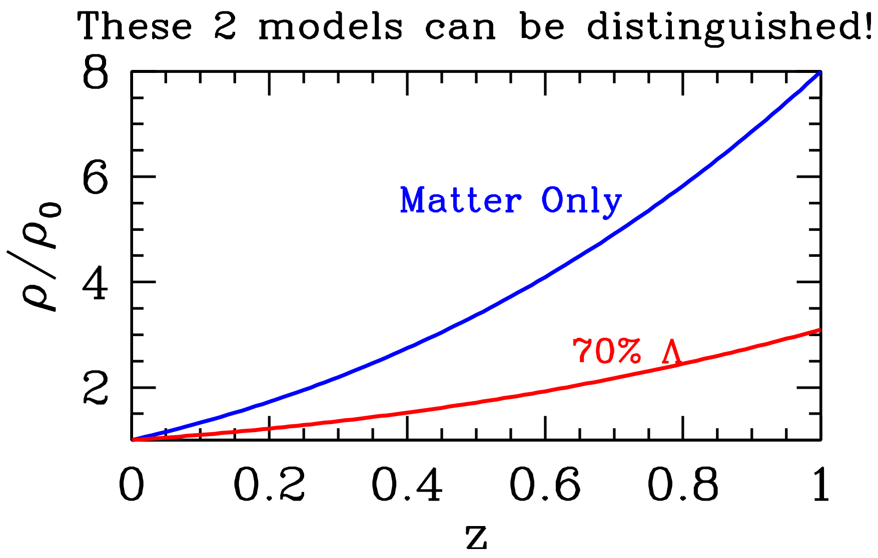
Evidence for Dark Energy

$$\rho \propto a^{3(1+w)} = (1+z)^{-3(1+w)}$$

- Matter density scales as a^{-3} $(w = 0)$
- Radiation scales as a^{-4} $(w = 1/3)$
- Cosmological constant is . . . constant $(w = -1)$



Evidence for Dark Energy



Expansion rate was slower in Λ model ↔ The universe is accelerating!

What observables depend on $H(z)$?

- Age of the universe:

$$t = \int_0^\infty \frac{dz}{H(z)(1+z)}.$$

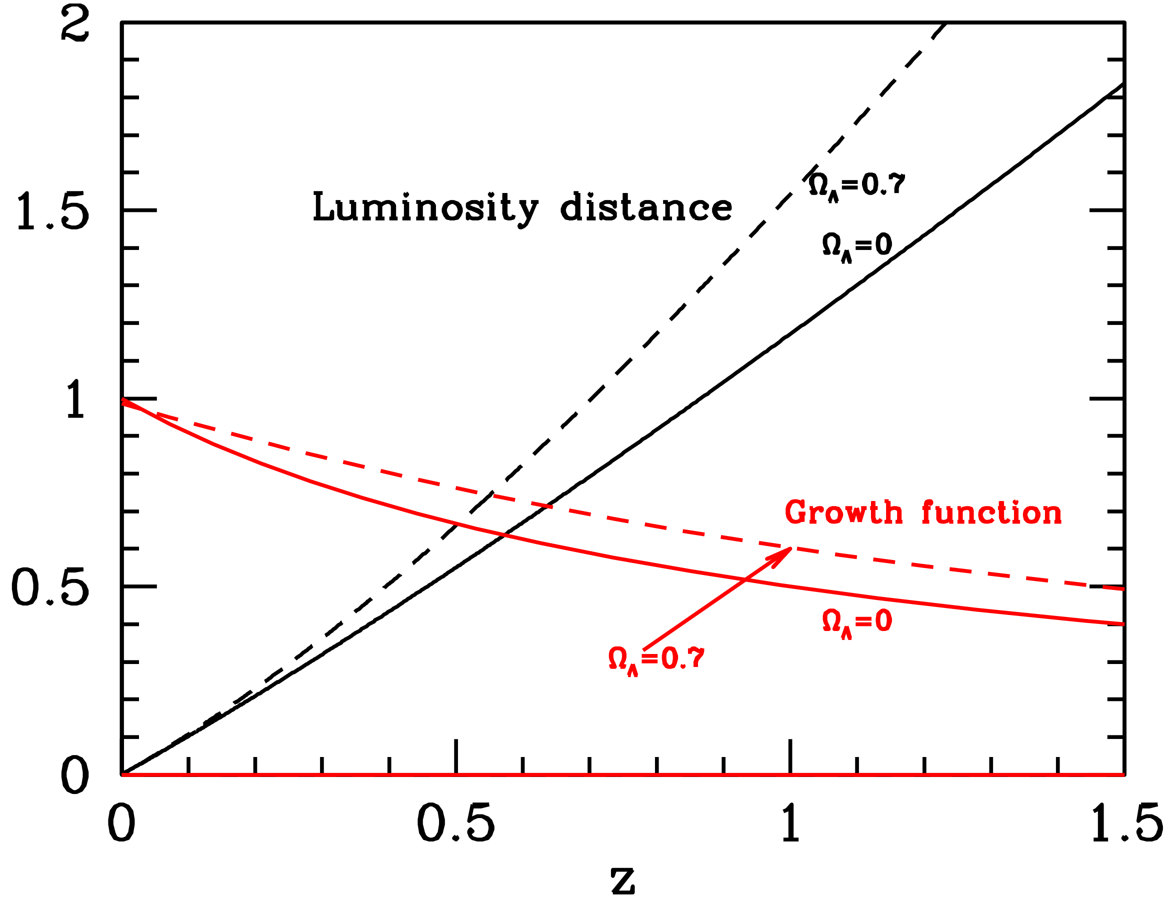
- Luminosity distance:

$$d_L(z) = (1+z)\int_0^z \frac{dz'}{H(z')}$$

- Growth function:

$$D_1(z) = \frac{5\Omega_m}{2}\frac{H(z)}{H_0}\int_0^z \frac{dz'\,(1+z')}{(H(z')/H_0)^3}$$

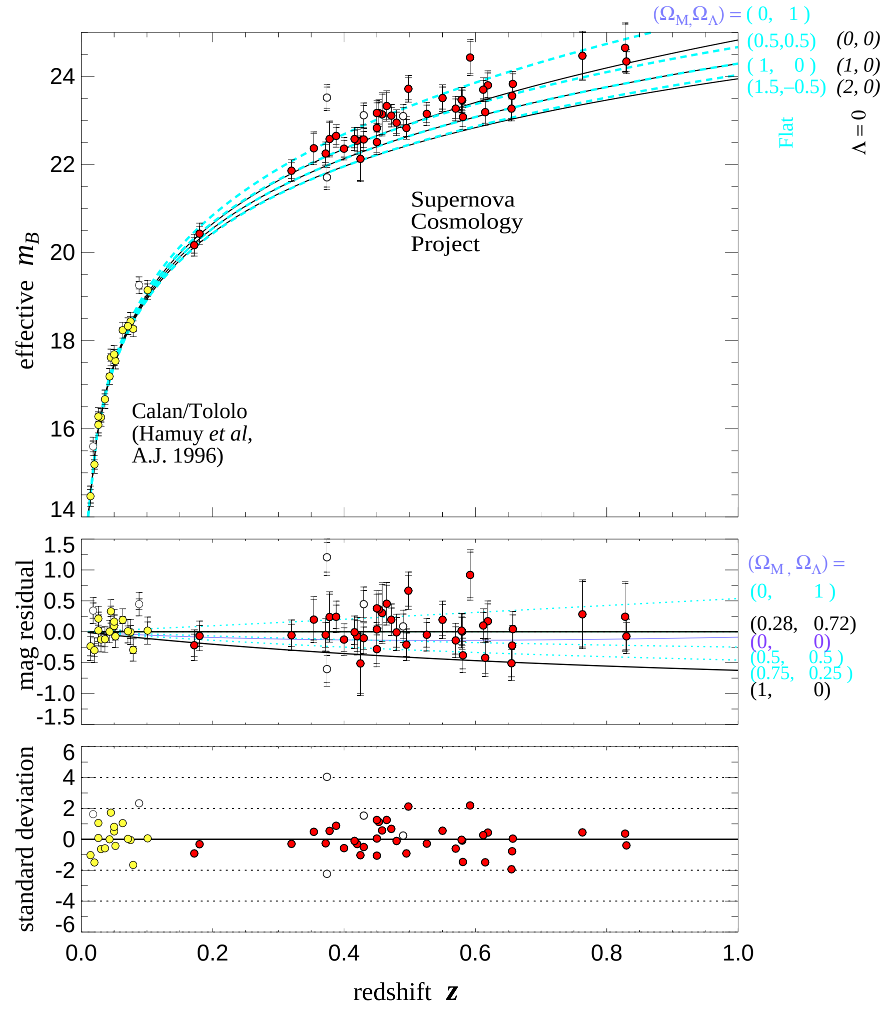
Evidence for Dark Energy



Evidence for Dark Energy: d_L

Type Ia Supernovae

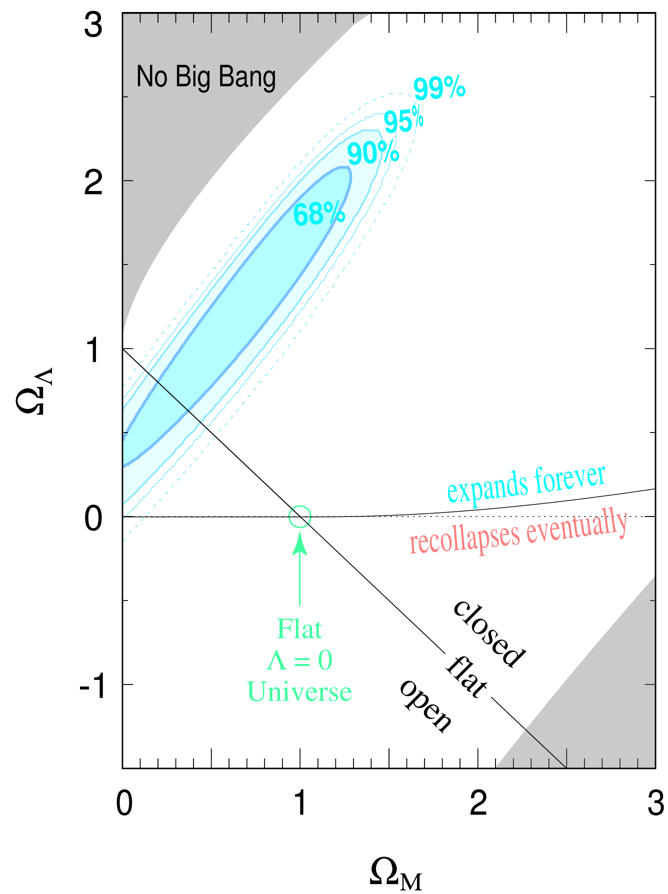
- Observed flux proportional to L/d_L^2. Type Ia SN are *standard candles* (identical L), so their apparent magnitude is a measure of d_L
- $H(z)$ smaller in Λ model $\longrightarrow$ d_L larger $\longrightarrow$ fainter SN



Perlmutter, *et al.* (1998)

Evidence for Dark Energy: d_L

$>$ 50 SN observed by 2 teams imply $\Omega_\Lambda \neq 0$.

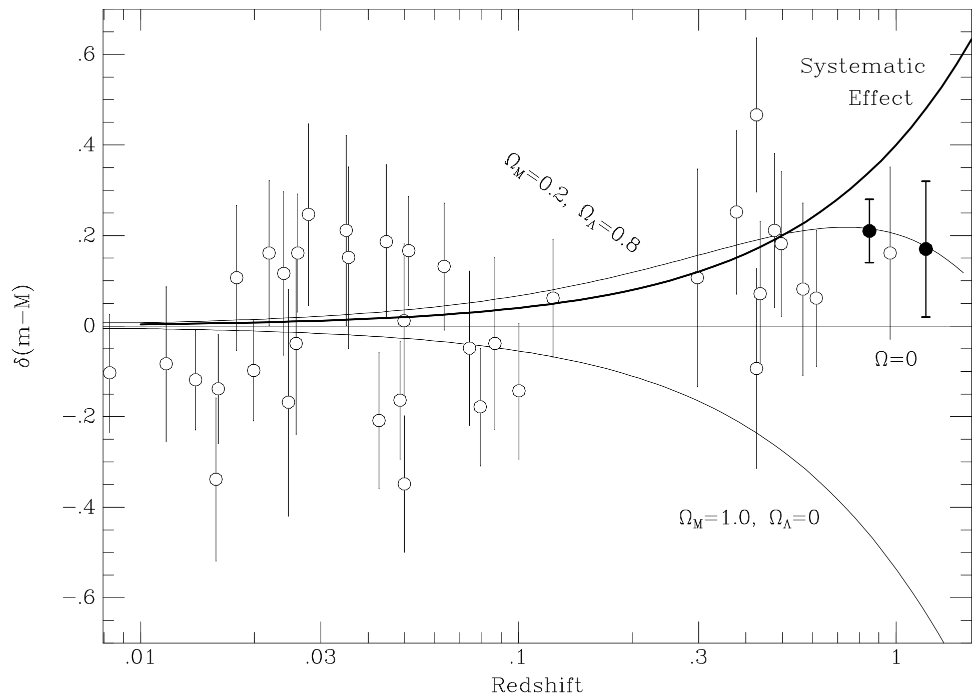


Riess et al. 1999

Evidence for Dark Energy: d_L

Systematic Effect?

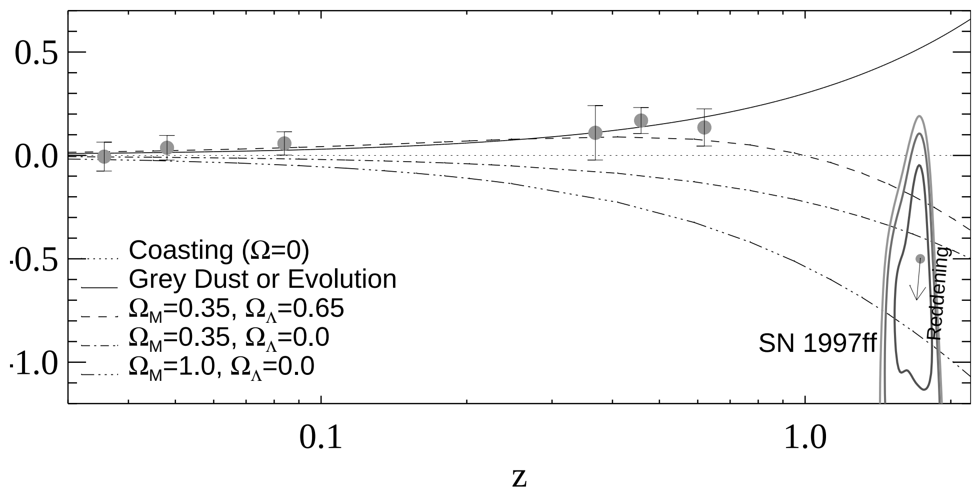
Ordinary dust *reddens* the image; this is not seen. Gray dust leads to lower fluxes as z increases



Riess et al. 1999

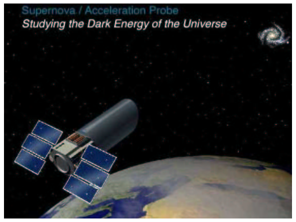
Evidence for Dark Energy: d_L

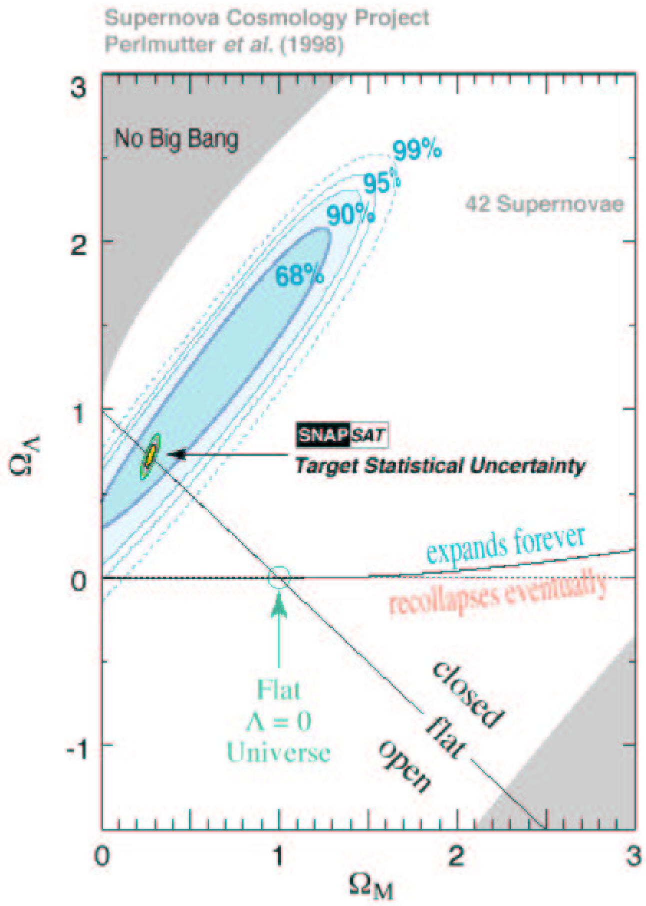
Recently SN observed at $z = 1.7$



Riess et al. 2001

Evidence for Dark Energy: d_L





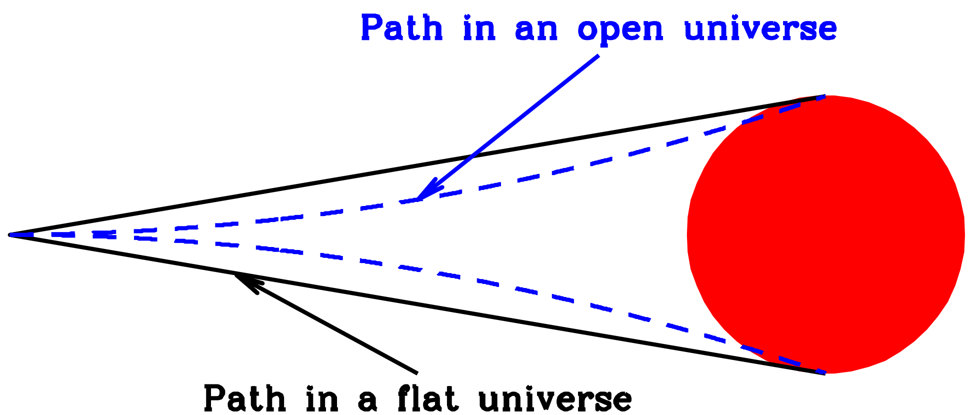
Proposed mission SNAP will observe thousands of distant SN at $z \simeq 1$

Evidence for Dark Energy: Inventory

- Hot/cold spots in CMB at $z = 1100$ are the size of sound horizon. Apparent size depends on geometry of universe.

- There are many estimates of matter density: all yield $\Omega_m = 0.3$.

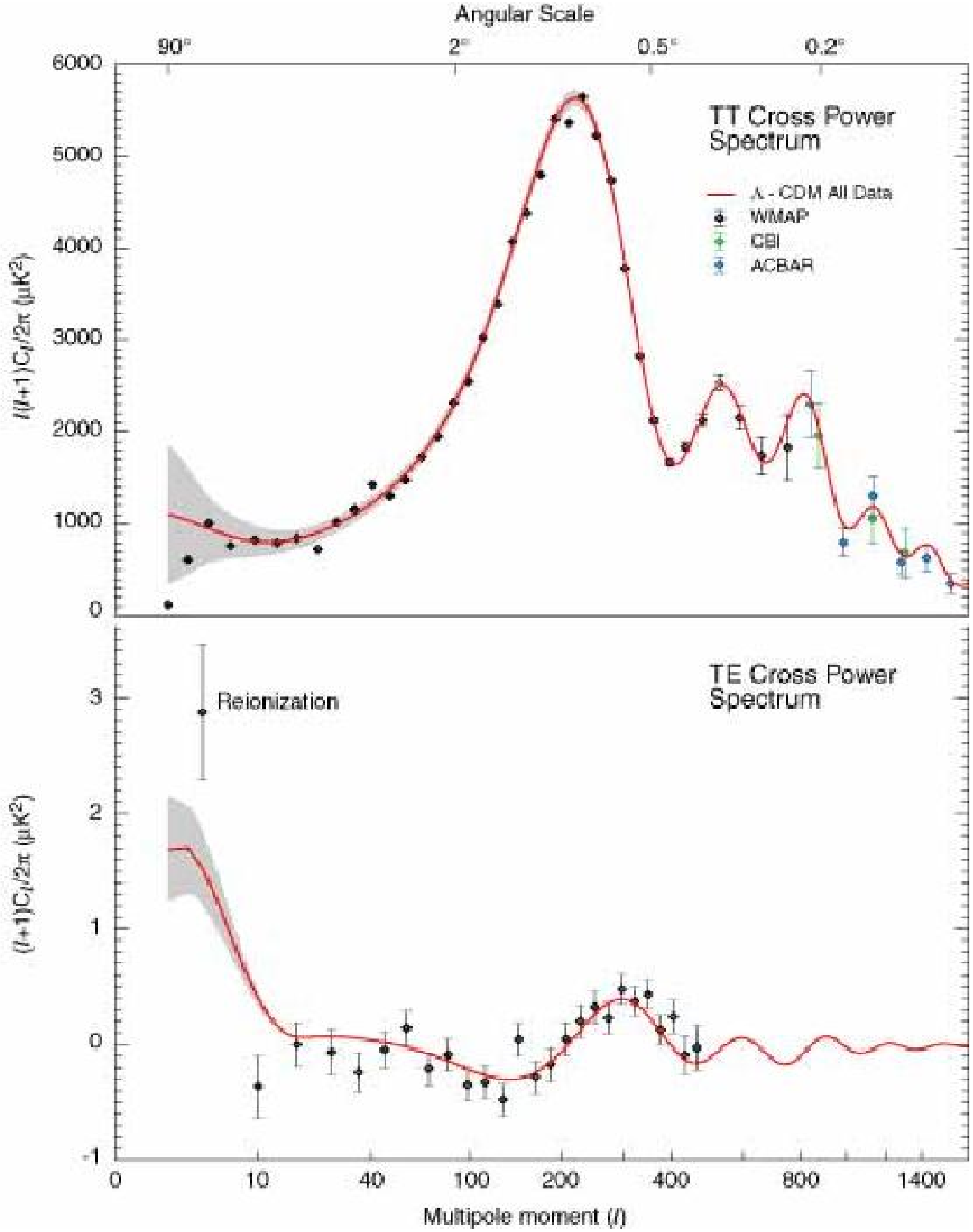
Evidence for Dark Energy: CMB



Angular size of hot/cold spots distinguishes between open, closed or flat universe.

Evidence for Dark Energy: CMB

Boomerang
DASI
Maxima
TOCO
CBI
Python
MSAM
Viper VSA
QMAP

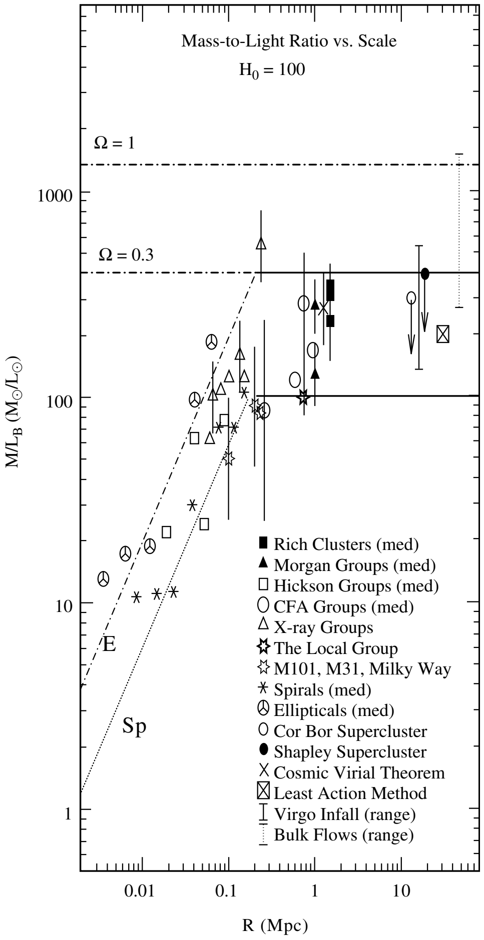


~ 10 experiments have verified position of first peak. Our universe is flat

$\longrightarrow$ Total energy density is equal to the critical density.

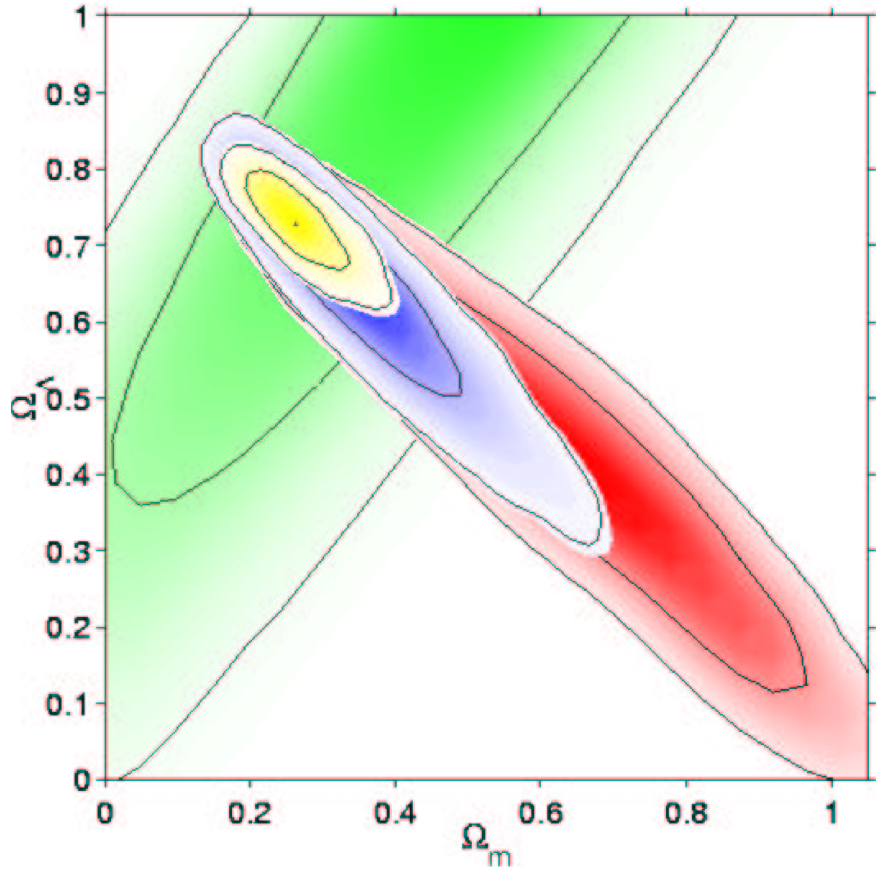
Evidence for Dark Energy: Ω_m

- Direct counting gives $\Omega_m = 0.3$
- Also: CMB, large scale structure, velocities . . . All give $\Omega_m = 0.3$



Evidence for Dark Energy

All data agree



SN

CMB

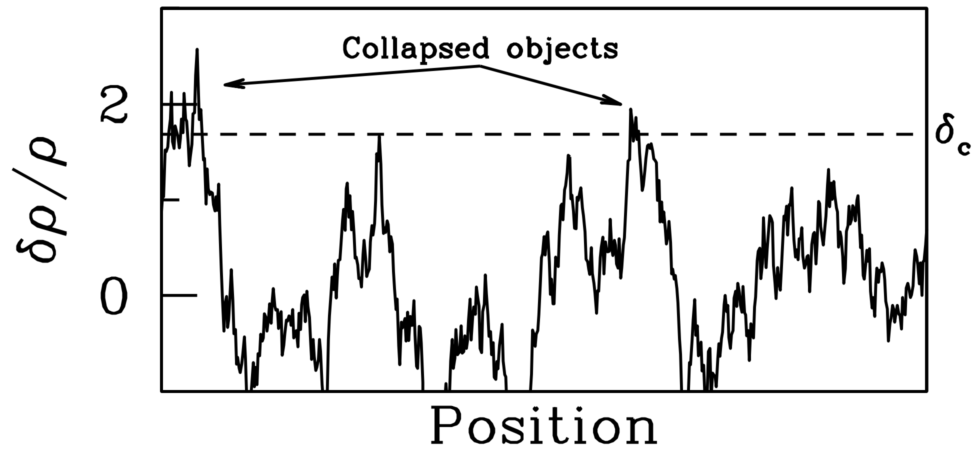
CMB+HST

ALL

Lewis & Bridle 2002

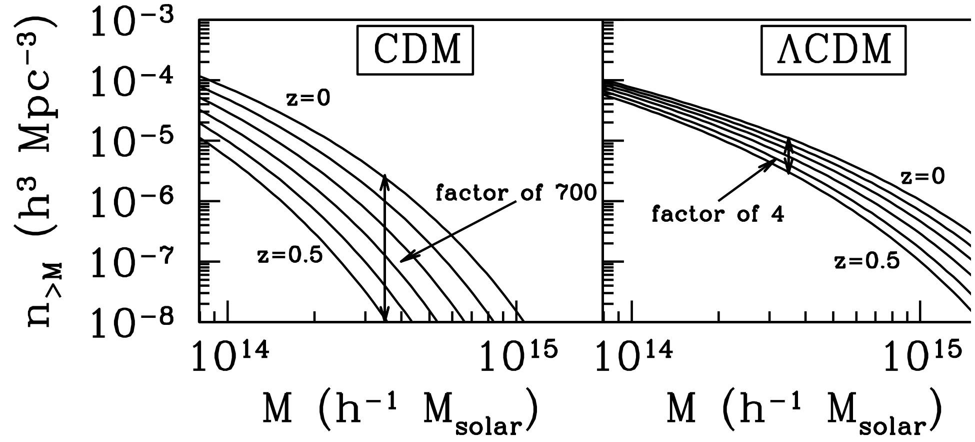
Evidence for Dark Energy: Growth function

- Large objects are rare: $\propto e^{-\delta^2/2\sigma^2}$
- High-sigma peaks in density distribution
- Exponentially sensitive to variance σ/growth function



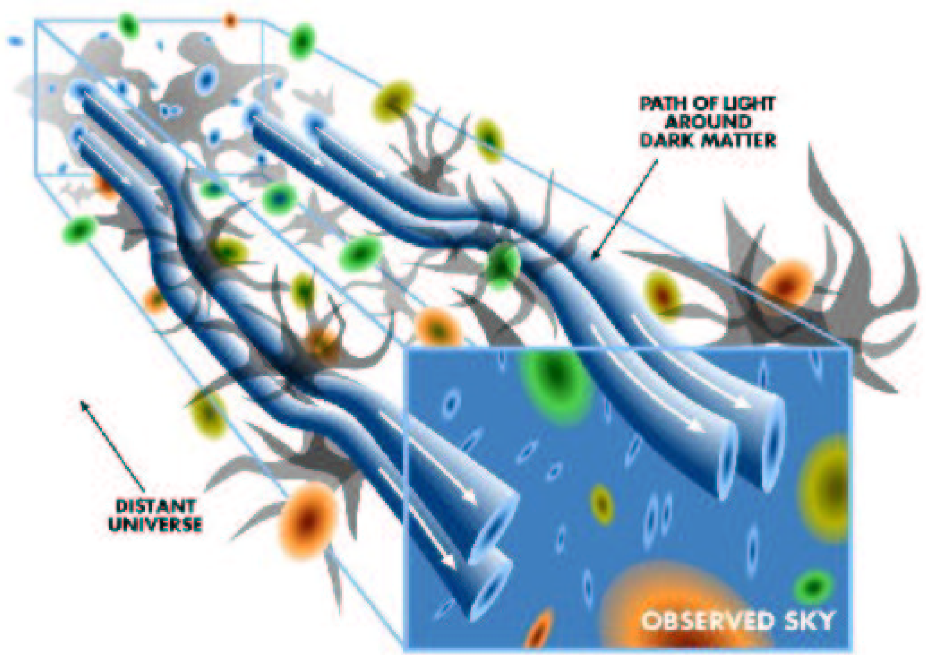
Evidence for Dark Energy: Growth function

- Less growth in a Λ universe
- Clustering was comparable at $z \sim 0.5 - 1$ to now
- Roughly same number of clusters



Evidence for Dark Energy: Growth function

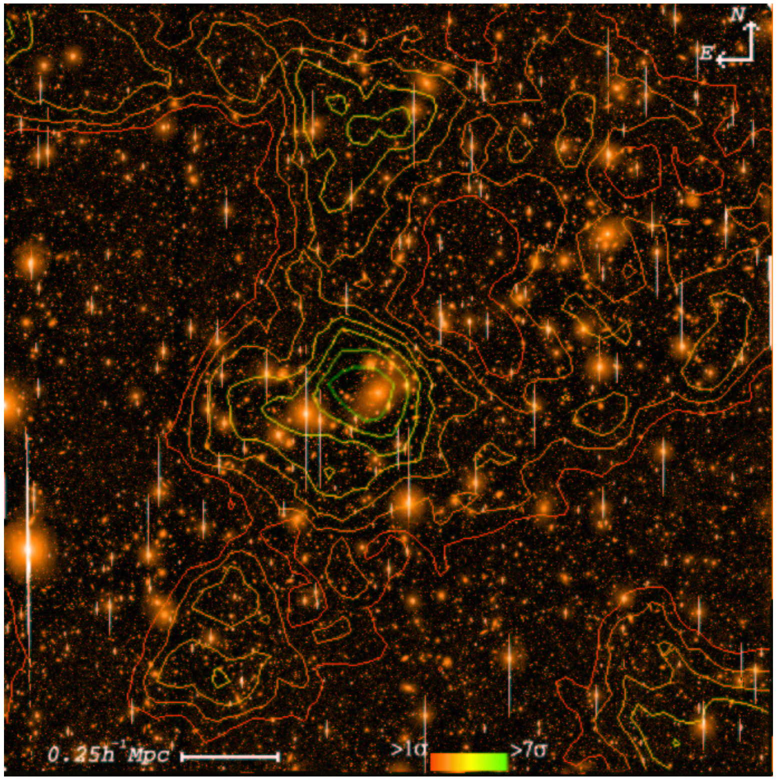
How can we measure mass?



Gravitational Lensing!

Evidence for Dark Energy: Growth function

Good way to measure cluster mass

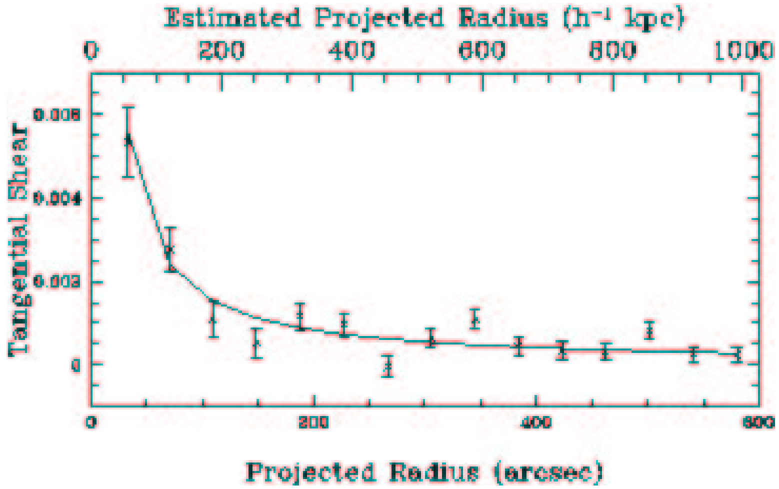


Joffre et al. 2000

Evidence for Dark Energy: Growth function

What can be done with lensing?

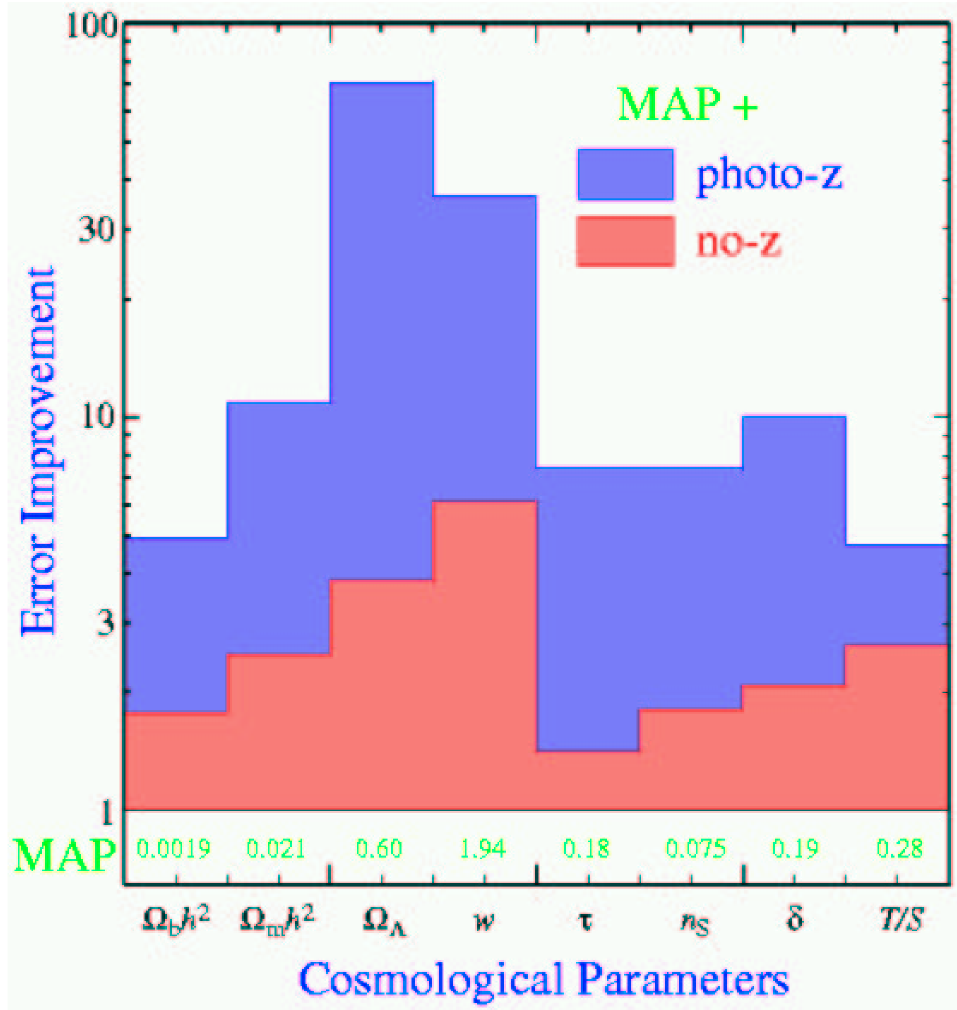
- cluster masses
- galaxy-galaxy
- lensing by lss
- lensing of cmb



SDSS: Fischer et al. 2000

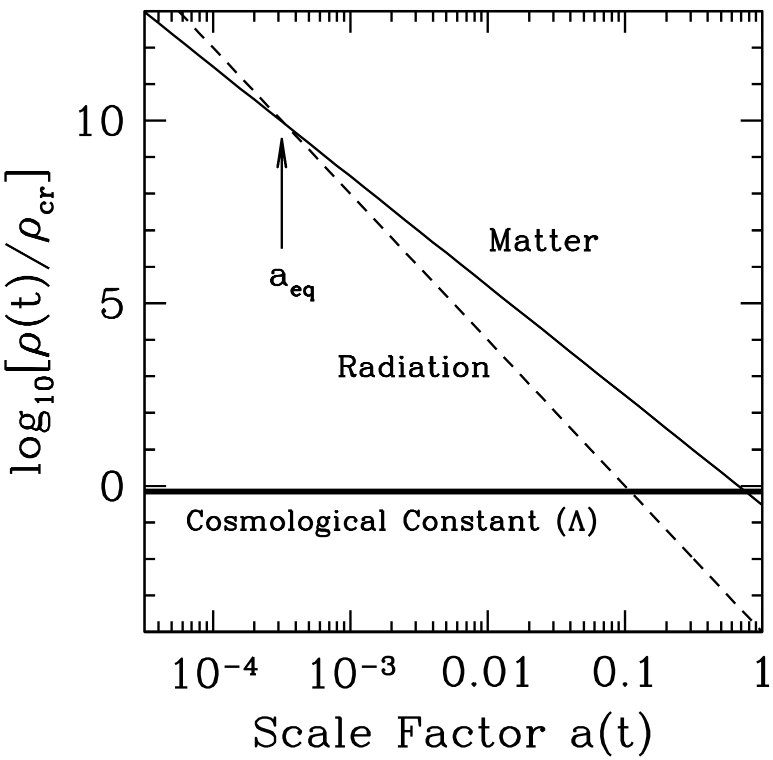
Evidence for Dark Energy: Growth function

Proposed LSST will measure weak lensing by large scale structure



What is it?

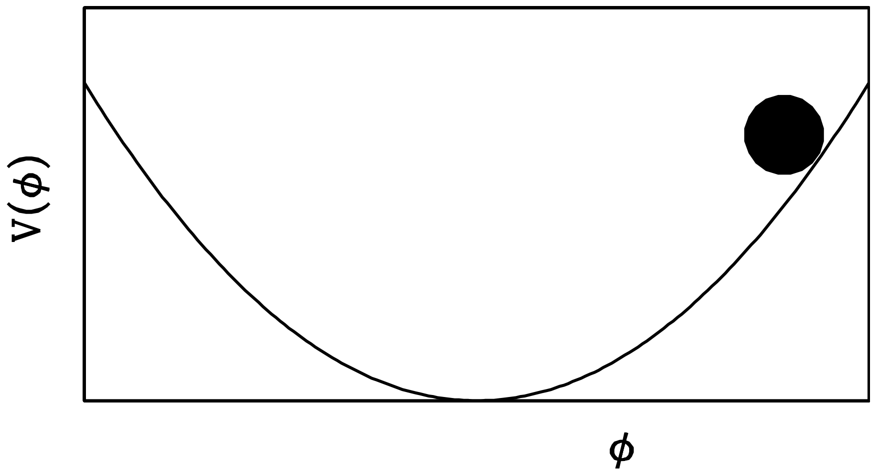
Why now? Now is the only time when $\rho_\Lambda \simeq$ ambient density. Need fine tuning initially to one part in 10^{128} to get present value.



What is it?

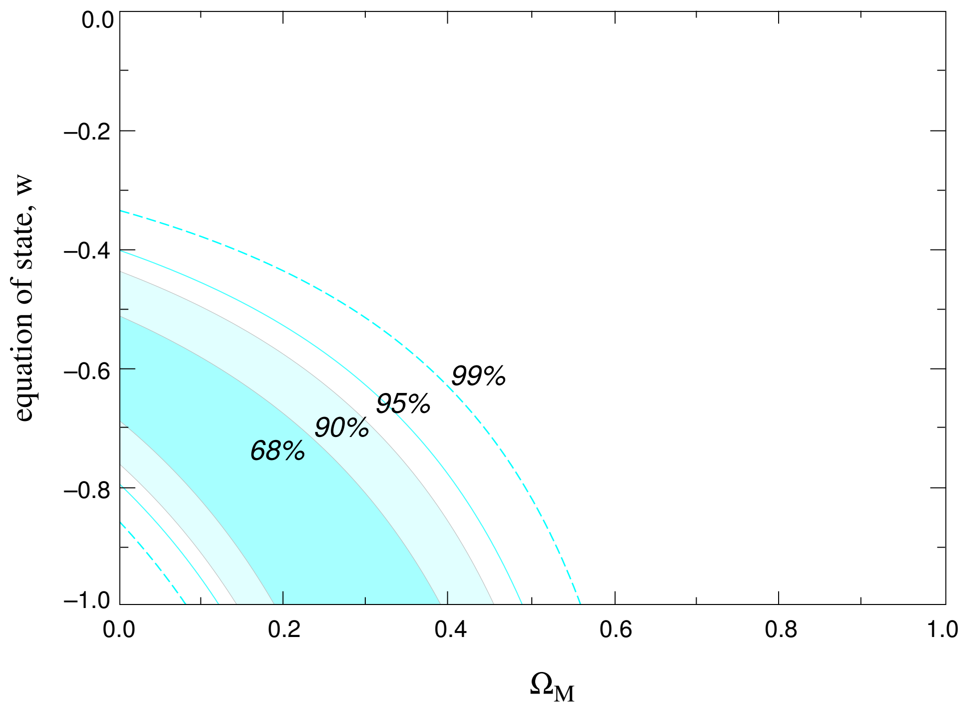
Quintessence

- True value of $\Lambda = 0$
- Some other form of energy non-zero today, will eventually relax to true vacuum.
- Most popular incarnation: single scalar field with $V(\phi)$



What is it?

Supernovae constraints can be re-expressed with w as free parameter. Require $w < -.5$; consistent with cosmological constant.

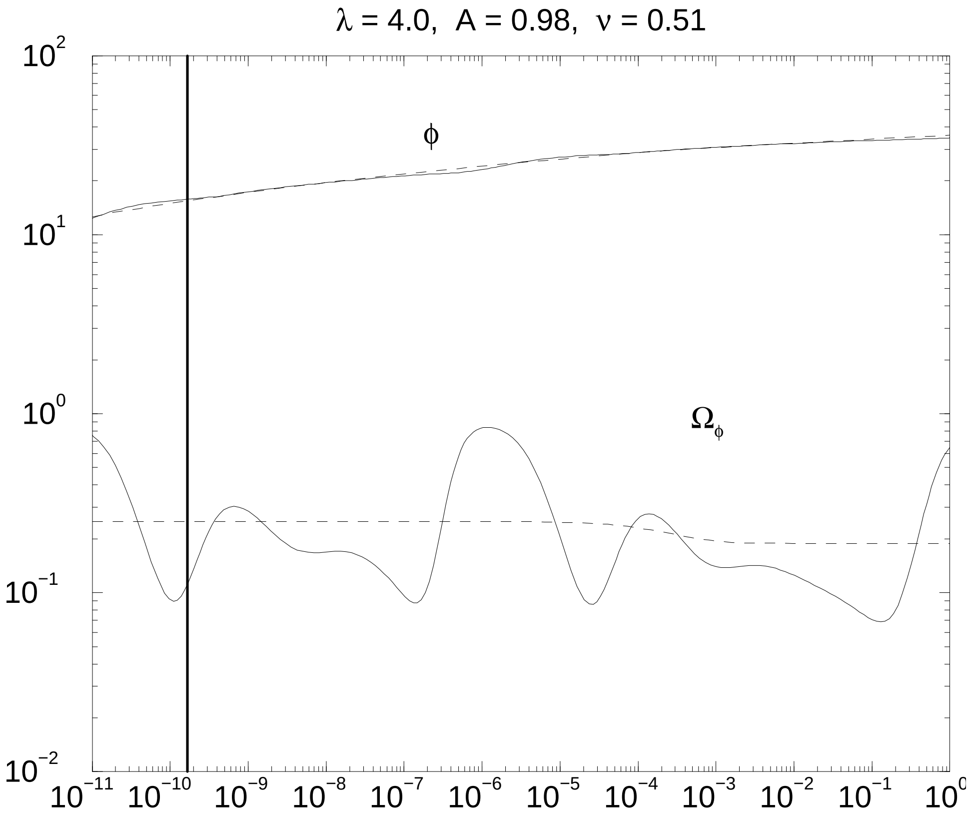


Riess et al. 1999

What is it?

My (Favorite) Quintessence Model

Exponential potential leads to ρ_ϕ tracking ambient density. Use instead $V(\phi) = e^{-\lambda\phi}\,[1 + A\sin(\nu\phi)]$



Dodelson, Stewart, & Kaplinghat (2001)

Conclusions

- Several pieces of independent evidence for dark energy: d_L, Type Ia supernovae and Cosmic Inventory. Efforts to hunt down systematics and increase statistics are ongoing
- Another class of evidence growth function (gravitational lensing, clusters) will play a key role in near future
- Apparently have discovered new form of energy which dominates the universe

MODERN COSMOLOGY

Scott Dodelson

Modern Cosmology begins with an introduction to the smooth, homogeneous universe described by a Friedmann-Robertson-Walker metric, including careful treatments of dark energy, big bang nucleosynthesis, recombination, and dark matter. From this starting point, the reader is introduced to perturbations about an FRW universe: their evolution with the Einstein-Boltzmann equations, their generation by primordial inflation, and their observational consequences. These consequences include the anisotropy spectrum of the cosmic microwave background (CMB) featuring acoustic peaks and polarization, the matter power spectrum with baryonic wiggles, and their detection via photometric galaxy surveys, redshift distortions, cluster abundances, and weak lensing. The book concludes with a long chapter on data analysis. *Modern Cosmology* is the first book to explain in detail the structure of the acoustic peaks in the CMB, the E/B decomposition in polarization which may allow for detection of primordial gravity waves, and the modern analysis techniques used on increasingly large cosmological data sets. Readers will gain the tools needed to work in cosmology and will learn how modern observations are rapidly revolutionizing our picture of the universe.

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